Depletion and low gas temperature in the L183 (= L134N) prestellar core: the N$_2$H$^+$ - N$_2$D$^+$ tool*

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ABSTRACT

Context. The study of pre-stellar cores (PSCs) suffers from a lack of undepleted species to trace the physical properties of the gas in their very dense inner parts.
Aims. We carry out detailed modelling of N$_2$H$^+$ and N$_2$D$^+$ cuts across the L183 main core to evaluate the depletion of these species and their usefulness as a probe of physical conditions in PSCs.
Methods. We have developed a non-LTE (NLTE) Monte-Carlo code treating the 1D radiative transfer of both N$_2$H$^+$ and N$_2$D$^+$, making use of recently published collisional coefficients with He between individual hyperfine levels. The code includes line overlap between hyperfine transitions. An extensive set of core models is calculated and compared with observations. Special attention is paid to the issue of source coupling to the antenna beam.
Results. The best-fitting models indicate that i) gas in the core center is very cold (7± 1 K) and thermalized with dust, ii) depletion of N$_2$H$^+$ does occur, starting at densities 5–7 $\times 10^5$ cm^{-3} and reaching a factor of 6^{+13}_{-3} in abundance, iii) deuterium fractionation reaches ~70% at the core center, and iv) the density profile is proportional to r^{-1} out to ~4000 AU, and to r^{-2} beyond.
Conclusions. Our NLTE code could be used to (re-)interpret recent and upcoming observations of N$_2$H$^+$ and N$_2$D$^+$ in many pre-stellar cores of interest, to obtain better temperature and abundance profiles.

Key words. ISM: abundances – ISM: molecules – Radiative transfer – ISM: structure – ISM: individual: L183 – Line: formation

1. Introduction

Understanding star formation is critically dependent upon the characterization of the initial conditions of gravitational collapse, which remain poorly known. It is therefore of prime importance to study the properties of pre-collapse objects, the so-called pre-stellar cores (hereafter PSCs). As bolometers and infrared extinction maps have been unveiling PSCs through their dust component, it has become clear that depletion of molecules onto ice mantles is taking place inside these cores, preventing their being studied with the usual spectroscopic tools. In most PSCs, very few observable species seem to survive in the gas phase in the dense and cold inner parts, namely N$_2$H$^+$, NH$_3$, H$_2$D$^+$, and their isotopologues (e.g. Tafalla et al. 2002). In a few cases, it is advocated that even N-bearing species also deplete (e.g. B68: Bergin et al. 2002; L1544: Walmsley et al. 2004; L183: Pagani et al. 2005, hereafter PPABC).

Among the above three species, H$_2$D$^+$ is thus the only one not to deplete. However, it has only one (ortho) transition observable from the ground, that moreover requires excellent sky conditions. The para form ground transition at 1.4 THz should not be detectable in emission in cores with T$_{kin}$ ≤ 10 K. Therefore H$_2$D$^+$ is useful for chemical and dynamical studies, but brings little information on gas physical conditions. The NH$_3$ inver-

sion lines at 23 GHz can provide kinetic temperature measurements as long as higher lying non-metastable levels are not significantly populated (Walsmley & Ungerechts 1983). However, because the critical density of the NH$_3$ (1,1) inversion line is only ~2000 cm^{-3}, this tracer may have a substantial contribution from external, warmer layers, not representative of the densest parts of PSCs. The third species, N$_2$H$^+$, appears very promising, because it has the strong advantage of having mm transitions with critical densities in the range 0.5–70 $\times 10^6$ cm^{-3} and intense hyperfine components for the (J:1–0) line in typical PSC conditions. The ratio of hyperfine components gives an estimate of the opacity and excitation temperature of the line. Fitting both the hyperfine ratios and the (J:1–0) to (J:3–2) ratio with a rigorous NLTE model should thus bring strong constraints on both T$_{kin}$ and n(H$_2$). Collisional coefficients (with He) between individual hyperfine sublevels have become available recently (Daniel et al. 2005). However, current excitation models (Daniel et al. 2006a) do not take line overlap into account, which limits their accuracy. Another question that remains is whether N$_2$H$^+$ is indeed able to probe the central core regions, despite the depletion effects that have been reported.

In this paper, we introduce a new NLTE Monte-Carlo 1D code[1] treating N$_2$H$^+$ and N$_2$D$^+$ radiative transfer with line overlap, and apply it to detailed analysis of the main PSC in L183, a clear-cut case of N$_2$H$^+$ depletion (cf. PPABC). We demonstrate the capability of the model to constrain physical conditions inside the PSC (temperature, density profile, abundance and deple-

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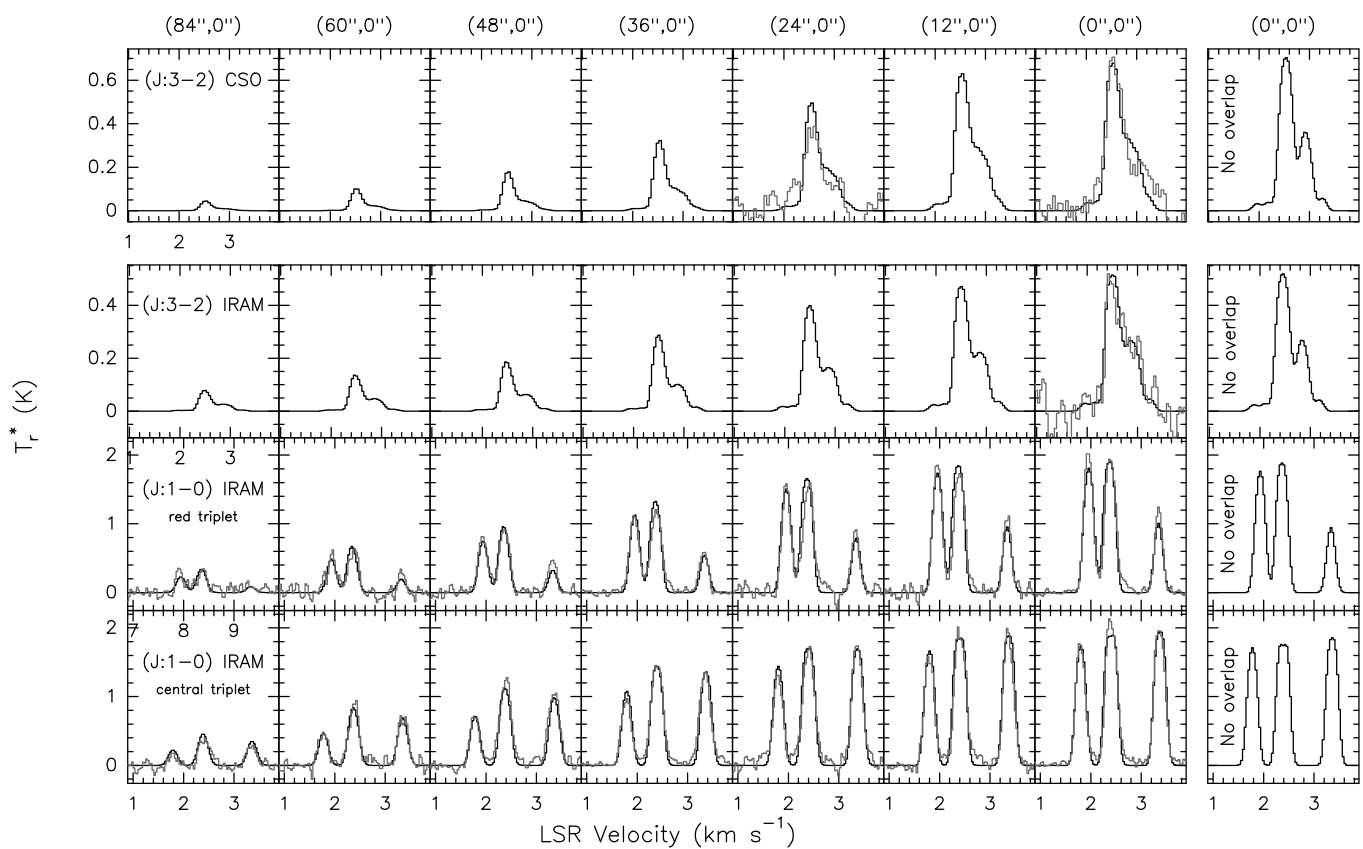[1] The Fortran code is available upon request to the author

Fig. 1. N_2H^+ observations (grey lines) compared to our best radiative transfer model (see Fig 4a-b). Each column corresponds to a spatial offset (indicated above the top box) from the main dust peak. The top row contains the CSO 10-m data, the other rows the IRAM 30-m data. The (J:1–0) line is split between the bottom row (central triplet) and the next row ('red' triplet); the isolated 'blue' component is not shown. The rightmost column represents the fit for the central position with a NLTE code without overlap for the same cloud parameters.

tion, deuterium fractionation). In particular we show that the gas is very cold at the core center and thermalized with the dust, and that N_2D^+ appears a very useful tracer of physical conditions in the innermost core regions.

2. Observations and analysis

The main dust core in L183 (Pagani et al. 2004) was observed at the IRAM 30-m telescope in November 2003, July 2004, and October 2006. Spectra were taken on a $12''$ grid in an East-West strip across the core (centered at $\alpha(2000) = $ 15h54m08.5s, $\delta(2000) = -2°52'48''$). Symmetric eastern and western spectra were averaged out to $\pm 48''$ to give a more representative radial profile. A small anti-symmetric velocity shift of a few tens of m s^{-1} was noticed on either side of the PSC center at distances beyond $\pm 30''$, possibly indicative of rotation (see Section 3.2). This shift was compensated for when averaging eastern and western spectra, in order not to artificially broaden the lines. Beyond $48''$, only eastern positions are considered, as the western side is contaminated by a separate core (Peak 3; Pagani et al. 2004). Lines of N_2H^+ (J:1–0), (J:3–2) and of N_2D^+ (J:1–0), (J:2–1) and (J:3–2) were observed in frequency-switching mode, with velocity sampling 30–50 m s^{-1} and T_{sys} ranging from 100 K at 3mm up to 1000 K at 1mm. The N_2H^+ (J:2–1) line at 186 GHz was not observed, as it lies only 3 GHz away from the telluric water line at 183 GHz, hence its usefulness for constraining excitation conditions would be limited by calibration sensitivity to even tiny sky fluctuations. The problem does not apply to N_2D^+, which lies a factor of 1.2 lower in frequency. Spatial res-

olution ranges from $33''$ at 77 GHz to $9''$ at 279 GHz. Additional CSO 10-m observations of N_2H^+ (J:3–2) were obtained in June 2004 at selected positions. Observations were done in position-switch mode with the high-resolution AOS (48 kHz sampling) and a T_{sys} around 600 K. A major interest in observing the (J:3–2) line with the CSO is that the beam size and efficiency are very similar to the 30-m values for the (J:1–0) line, thus almost canceling-out beam correction errors in the comparison between the two. It is thus a useful constraint for radiative transfer modeling, even though the higher-resolution ($9''$) 30-m (J:3–2) observations will be essential in constraining abundances and physical conditions in the innermost core regions. Standard reduction techniques were applied using the CLASS reduction package[2].

Complementary observations of NH_3 (1,1) and (2,2) inversion lines towards the PSC center were also obtained at the new Green Bank 100-m telescope (GBT) in November 2006 with velocity sampling of 20 m s^{-1} and a typical T_{sys} of 50 K, in frequency-switch mode. The angular resolution ($\sim35''$) is close to that of the 30-m in the low-frequency (J:1-0) N_2D^+ line. The main beam efficiency of the GBT at 23 GHz is between 0.95 and 1, so no correction was applied to put the spectra on T_R^* scale.

Table 1 summarizes the main observational parameters of the lines observed towards the core center: noise level of the spectrum (rms), rotational line total opacity, and intrinsic FWHM width of individual hyperfine transitions (as fitted by the CLASS HFS routine, which assumes equal excitation temperature for all sublevels), and relative velocity centroids and intensities of the main hyperfine groups in our spectra (derived from Gaussian

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Table 1. Spectral parameters of each line observed towards the PSC center position[a]

	$N_2H^+(J:1-0)$					$N_2H^+(J:3-2)$(IRAM)			$N_2H^+(J:3-2)$(CSO)		
Noise (1σ) :	28 mK					59 mK			64 mK		
Ref. frequency :	93.173764 GHz					279.511832 GHz			279.511832 GHz		
Intrinsic FWHM[b] :	0.195(1) km s^{-1}					0.18(2) km s^{-1}			0.28(4) km s^{-1}		
Total opacity[b] :	21.8(5)					10(2)			2(1)		
Hyperfine Group[c,d]	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})	T^*_R ratio[e] obs	weights	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})
1	−8.009(7)	2.05(3)	0.47(3)	0.932	*0.428*	−2.58(1)	0.17(6)	0.013(5)	−2.63(2)	0.20(6)	0.025(8)
2	−0.611(8)	1.81(3)	0.42(3)	0.822	*0.428*	0.00(1)	0.50(6)	0.29(2)	0.00(1)	0.67(6)	0.37(2)
3	0.000(7)	2.20(3)	0.60(3)	1	*1*	4.70(8)	2.25(6)	0.014(5)	4.68(3)	2.12(6)	0.028(9)
4	0.956(7)	2.08(3)	0.54(3)	0.945	*0.714*						
5	5.546(7)	2.12(3)	0.47(3)	0.964	*0.428*						
6	5.983(8)	2.04(3)	0.53(3)	0.927	*0.714*						
7	6.94(1)	1.26(3)	0.24(3)	0.573	*0.143*						

	$N_2D^+(J:1-0)$					$N_2D^+(J:2-1)$			$N_2D^+(J:3-2)$		
Noise (1σ) :	31 mK					26 mK			43 mK		
Ref. frequency :	77.109616 GHz					154.217182 GHz			231.321917 GHz		
Intrinsic FWHM[b] :	0.226 (3) km s^{-1}					0.185 (2) km s^{-1}			0.18 (2) km s^{-1}		
Total opacity[b] :	4.7(3)					4.9(1)			1.5(7)		
Hyperfine Group[c,d]	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})	T^*_R ratio[e] obs	weights	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})
1	−9.697(8)	0.77(3)	0.15(1)	0.7	*0.428*	−5.35(2)	0.16(3)	0.137(5)	−3.26(1)	< 0.12	
2	−0.763(8)	0.76(3)	0.15(1)	0.691	*0.428*	−0.759(8)	0.21(3)	0.080(3)	0.000(5)	0.72(4)	0.186(8)
3	0.000(7)	1.10(3)	0.33(2)	1	*1*	0.000(1)	1.37(3)	0.513(3)	0.46(2)	0.19(4)	0.06(1)
4	1.146(8)	0.92(3)	0.25(2)	0.836	*0.714*	0.646(4)	0.40(3)	0.139(3)	2.47(3)	0.11(4)	0.028(7)
5	6.65(2)	0.55(3)	0.16(2)	0.5	*0.428*	2.5(1)	0.08(3)	0.02(2)			
6	7.19(1)	0.90(3)	0.23(2)	0.818	*0.714*	2.97(2)	0.43(3)	0.08(2)			
7	8.34(2)	0.34(3)	0.05(2)	0.309	*0.143*	3.57(5)	0.26(3)	0.16(3)			

	$NH_3(1,1)$					$NH_3(2,2)$					
Noise (1σ) :	74 mK					76 mK					
Ref. frequency :	23.6944957 GHz					23.7226332 GHz					
Intrinsic FWHM[b] :	0.195 (1) km s^{-1}					0.20 (1) km s^{-1}					
Total opacity[b] :	24.2(4)					0.1(7)					
Hyperfine Group[c,d]	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})	T^*_R ratio[e] obs	weights	$V_{rel.}$ (km s^{-1})	T^*_R (K)	*Area* (K km s^{-1})			
1	−19.504(6)	2.62(7)	0.89(4)	0.922	*0.363*	−0.005(5)	0.73(8)	0.152(7)			
2	−7.814(5)	2.68(7)	0.69(3)	0.944	*0.273*						
3	−7.255(8)	1.95(7)	0.57(3)	0.687	*0.182*						
4	−0.17(1)	2.84(7)	1.15(9)	1	*1*						
5	0.30(2)	2.78(7)	1.0(1)	0.979	*0.636*						
6	7.465(9)	2.60(7)	0.72(5)	0.915	*0.303*						
7	7.89(1)	1.91(7)	0.48(5)	0.673	*0.152*						
8	19.318(9)	2.48(7)	0.62(4)	0.873	*0.242*						
9	19.85(1)	1.72(7)	0.42(5)	0.606	*0.121*						

[a] $\alpha(2000)$ = 15h54m08.5s, $\delta(2000)$= −2°52′48″ with V_{LSR} = 2.3672(2) km s^{-1} (from the NH_3 (1,1) measurement). Values in parentheses represent the uncertainty (1σ) on the last digit of the tabulated values.

[b] Intrinsic linewidth of hyperfine transitions and total opacity from the CLASS HFS routine. Note that for CSO observations are done with a resolution around 0.1 km s^{-1} (or more) which explains this larger value

[c] Relative velocity and intensity of the main hyperfine groups (from gaussian fits).

[d] The detailed hyperfine structure of the 3 species is described in Caselli et al. (1995) for N_2H^+, Dore et al. (2004) for N_2D^+ and Kukolich (1967) for NH_3

[e] For comparison, the T^*_R ratios of hyperfine groups given by their statistical weights are listed in italics.

fits). Note that the hyperfine groups are always broader than the intrinsic linewidth derived by CLASS, due to optical-depth effects and/or to the presence of adjacent hyperfine components too close to be spectroscopically separated.

To compare with models, beam efficiency correction is an important issue. Since L183 (as well as other PSCs) is very extended compared to the primary beam of 10″–30″, the T_{mb} intensity scale is inadequate, as it will overcorrect for source-antenna coupling and thus overestimate the true source surface

brightness[3] (cf. PPABC). To avoid this problem, observed spectra are corrected only for moon efficiency ($\sim T_R^*$ scale), while models are convolved by the full beam pattern of the 30m telescope (cf. Greve et al. 1998)[4], yielding intensities on the same T_R^* scale. For the CSO, we also correct the data for moon efficiency ($\eta_{\mathrm{moon}} \approx 0.8$ at 279 GHz); and as the details of the beam pattern are not known, the model output is convolved with a simple Gaussian beam. Because the CSO main beam coupling is very good at 1.1 mm ($\eta_{MB} \geq 0.7$), the uncertainty of the correction should be limited.

To correctly treat the beam coupling to the source, we also take into account that the L183 core is located within a larger-scale N_2H^+ filament oriented roughly north-south, as revealed by previous low-resolution maps (PPABC). This elongation is mostly constant in intensity over $\simeq 6$ arcminutes. Therefore we approximate the brightness distribution as a cylinder of 6′length, with its axis in the plane of the sky. The intensity distribution along the east-west section of the cylinder is taken as the emergent signal along the equator of our spherical Monte-Carlo model. We then replicate these values along the (north-south) cylinder axis before convolving by the antenna beam. For N_2D^+, the (unpublished) large-scale map shows a smaller north-south extent ($\simeq 2'$), so we use a cylinder length of 2′ for convolving our model.

In Figs. 1 and 2, we plot the observations in T_R^* compared to our best model (see Sect. 3.4), convolved by the antenna beam as explained above.

3. Modeling and discussion

3.1. Radiative transfer code

Our spherical Monte-Carlo code is derived from the Bernes code (Bernes 1979) and has been extensively tested on models of CO emission from dense cores (e.g. Pagani 1998). It was recently updated to take into account overlap between hyperfine transitions occuring at close or identical frequencies. This is done by treating the photon transfer for all hyperfine transitions of the same rotational line simultaneously (see also Appendix in Gónzalez-Alfonso & Cernicharo 1993). Instead of randomly choosing the frequency of the emitted photon, a binned frequency vector is defined for each rotational transition inside which all hyperfine transitions are positioned. All bins are filled with their share of spontaneously emitted photons (or 2.7 K continuum background photons) and, during photon propagation, absorption is calculated for each bin by summing all the hyperfine transition opacities at that frequency

$$\tau(\nu) = \int \kappa(\nu)\mathrm{d}s = \sum_i \tau_i \phi_i(\nu) = \int \left(\sum_i \kappa_i \phi_i(\nu) \right) \mathrm{d}s$$

where κ_i is the absorption coefficient, τ_i the opacity and $\phi_i(\nu)$ the local frequency profile of the ith hyperfine transition. The total number of absorbed photons $I_o(\nu)e^{-\tau(\nu)}$ is then redistributed

among hyperfine levels according to their relative absorption coefficients

$$\mathrm{d}I_i(\nu) = \frac{\kappa_i \phi_i(\nu)}{\kappa(\nu)} \times I_o(\nu)e^{-\tau(\nu)}.$$

Overlap is treated for all rotational transitions. Statistical equilibrium, including collisions, is then solved separately for each hyperfine energy level.

Collisional coefficients are available for transitions between all individual hyperfine levels, but were calculated with He as a collisioner instead of H_2 (Daniel et al. 2005). We scaled them up by 1.37 to correct for the difference in reduced mass, but note that this correction is only approximate (see Willey et al. 2002 and Daniel et al. 2006a). The code works for both isotopologues. As one can probably neglect the variation in the electric dipole moment (Gerin et al. 2001), the N_2D^+ Einstein spontaneous emission coefficients are derived from those of N_2H^+ by simply scaling them down by 1.2^3. Deexcitation collisional coefficients are kept the same as for N_2H^+ (following Daniel et al. 2006b).

The local line profile in our 1D code takes both thermal and turbulent broadening into account, as well as any radial and rotational velocity fields. In principle, rotation introduces a deviation from spherical symmetry. However, for small rotational velocities (less than half the linewidth, typically), the excitation conditions do not change noticeably within a given radial shell, as shown by a previous 2–D version of this code (Pagani & Bréart de Boisanger 1996). Therefore the (much faster) 1D version remains sufficiently accurate.

Our Monte-Carlo model shows that excitation temperatures among individual hyperfine transitions within a given rotational line differ by up to 15 %; hence, the usual assumption of a single excitation temperature to estimate the line opacity (e.g. in the CLASS HFS routine) is not fully accurate as already noticed by Caselli et al. (1995). Neglecting line overlap affects the excitation temperature of hyperfine components differentially when the opacity is high enough, mostly decreasing it, but also increasing it in a few cases. For example, in our best model for the L183 PSC, the excitation temperatures of individual hyperfine components change by up to 10%. As shown in the last column of Fig. 1, a noticeable difference between models with and without overlap appears in the (J:3–2) line shape.

3.2. Grid of models

Our core model has an outer radius of 13333 AU $\simeq 2'$, fixed by the maximum width $\simeq 4'$of detectable N_2H^+ emission in the east-west direction across the PSC, as seen in large-scale maps (PPABC). We assume that N_2H^+ abundance is zero outside of this radius (due to chemical destruction by CO).

For all models, the microscopic turbulence is set to Δv_{turb}(FWHM) = 0.136 km s^{-1}. This was found sufficient to reproduce the width of individual hyperfine groups, and is comparable to the thermal width contribution. A small rotational velocity field of a few tens of m s^{-1} was also imposed beyond 3000 AU (see Table 2) to reproduce the small anti-symmetric velocity shift from PSC center observed at distances beyond $\pm$ 30″ (cf. Sect. 2). The radial velocity was kept to 0 in all layers. As seen in Figs. 1 and 2, this approximation already gives a remarkable fit to observed line profiles, and is therefore sufficient for deriving the overall temperature and abundance structure.

We considered a variety of density profiles: single power law profiles $\rho \propto r^{-1}$, $\rho \propto r^{-1.5}$, $\rho \propto r^{-2}$, as well as broken power laws with $\rho \propto r^{-1}$ out to 4000 AU followed by r^{-2}, or
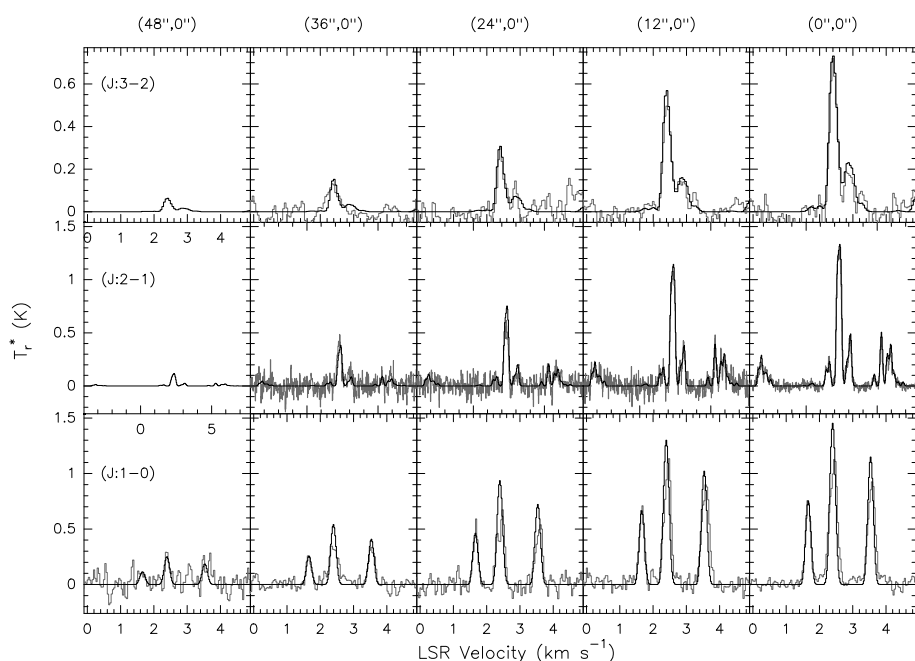
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[3] Teyssier et al. (2002) provide appropriate corrections for 30-m data, but only for uniform circular disk sources, and for data taken before the 1998 surface readjustment.

[4] The last surface readjustment of the 30-m in 1998 has not been characterized in detail, therefore we have scaled down the error-beam coupling coefficients given in Greve et al. (1998) to retrieve the new, improved beam efficiencies

Fig. 2. N$_2$D$^+$ observations (grey lines) compared to our best radiative transfer model for the largest N$_2$D$^+$ abundance possible (dotted histogram in Fig 4a). All data are from IRAM 30-m. Only the central triplet is represented for the (J:1–0) transition

by r$^{-3.5}$. For good accuracy, the density profile was sampled in 330 AU = 3″ thick shells, i.e. 3 to 10 times smaller than our observational beam sizes. In all cases, divergence at $r = 0$ was avoided by adopting inside r = 990 AU the density slope derived from the ISOCAM data (Pagani et al. 2004). The density profiles were then scaled to give the same total column density N(H$_2$) = 1.4 × 10^{23} cm^{-2} towards the PSC center, as estimated from the MAMBO emission map (Pagani et al. 2004). The effect of changing the total N(H$_2$) on the fit results is discussed in Sect. 3.4.1.

We then considered a variety of temperatures spanning the range 6–10 K in the inner core regions. We also tested non-constant temperature laws, increasing up to 12 K in the outermost layers and/or in the innermost layers.

For each combination of density and temperature laws, a range of N$_2$H$^+$ abundance profiles was investigated. To avoid a prohibitive number of cases, abundances were sampled only in 6 concentric layers, corresponding to the spatial sampling of our spectra (except the last, wider layer that encompasses the outermost two offsets), and multiple maxima/minima were forbidden. This led to a total of about 40,000 calculated models with 8 free parameters (6 abundances, 1 temperature profile, 1 density profile). Table 2 lists the values of H$_2$ densities for 4 density laws, together with the temperature and abundance profiles that best-fitted the data in each case (cf. next section).

3.3. Selection criteria

Our best-fit selection criteria are based on the use of several χ^2 and (reduced) χ_ν^2 estimates. Because the N$_2$H$^+$ (J:1–0) data have much more overall weight than the (J:3–2) data (more isolated hyperfine groups per spectrum, more observed positions, higher signal-to-noise ratio), computing a single χ_ν^2 per model is hardly sensitive to the quality of the (J:3–2) line fit. However, the (J:1–0) spectra being optically thick, they give few constraints on the temperature and N$_2$H$^+$ abundance of the innermost layers (within a 6″ radius). Those are only constrained by the IRAM (J:3–2) observations. The CSO (J:3–2) spectrum at (24″,0″) also adds interesting information on the density profile.
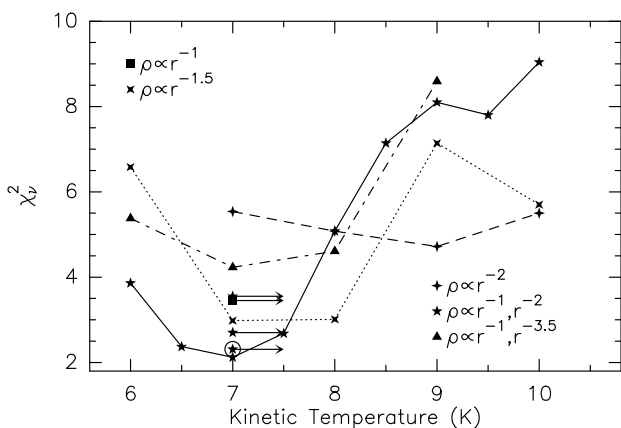

Fig. 3. Global χ_ν^2 for all density profiles and temperature cases. An arrow indicates that the majority of the model layers have the given temperature but that outer layers have rising temperatures. Only one (circled) solution satisfies all 4 individual χ^2 fittings (see text)

To constrain the models, we thus evaluate the goodness of fit independently for the (J:1–0) and (J:3–2) data, by computing χ^2 values on 4 types of measurements: 1) full area of each of the 7 (J:1–0) hyperfine components, at each of the 7 offset positions (49 values). This measurement is sensitive to the global temperature and abundance profiles. Fitting each hyperfine component separately allows solutions with overall identical area but different hyperfine line ratios to be discarded; 2) area of the 5 central channels (= 0.15 km s^{-1}) of each hyperfine component for the same lines (49 values), in order to reject self-absorbed profiles, i.e. line profiles that have the same total area but are wider and self-absorbed; 3) total area of the (J:3–2) CSO spectrum at (24″,0″) (1 value), a single measurement that helps to constrain the density profile; 4) total area of the (J:3-2) 30-m spectrum (1 value) that constrains the abundance (and the temperature to some extent) at the core center. These four measurements are of course not completely independent of each other, and trying to improve one fit by changing a model parameter can degrade one or several of the others. We used χ_ν^2 for the first two (with 41 degrees of freedom) and χ^2 for the last two (only one measure-

Table 2. The 4 main density laws investigated in our models, and the corresponding temperature and abundance profiles giving the best fit to N_2H^+ data in each case.

Radius[b] A.U.	V_{rot} (km s^{-1})	Density (cm^{-3})	T_{kin} (K)	Abundance N_2H^+	Abundance N_2D^+	Density (cm^{-3})	T_{kin} (K)	Abundance N_2H^+	Density (cm^{-3})	T_{kin} (K)	Abundance N_2H^+	Density (cm^{-3})	T_{kin} (K)	Abundance N_2H^+
		$\rho \propto r^{-1}, r^{-2}$ (Best model)[a]				$\rho \propto r^{-1}$			$\rho \propto r^{-1.5}$			$\rho \propto r^{-2}$		
330	0	2.34(6)[c,d]	7	2.4(−11)	2(−11)	2.34(6)[c]	7	7(−11)	3.30(6)[e]	7	1(−11)	3.46(6)[e]	9	1(−11)
660	0	2.05(6)[c]	7	2.4(−11)	2(−11)	2.05(6)[c]	7	7(−11)	2.89(6)[e]	7	1(−11)	3.03(6)[e]	9	1(−11)
990	0	1.55(6)[c]	7	8.5(−11)	4(−11)	1.55(6)[c]	7	7(−11)	2.19(6)[e]	7	7(−11)	2.29(6)[e]	9	3(−11)
1320	0	1.16(6)	7	8.5(−11)	4(−11)	1.16(6)	7	7(−11)	1.19(6)	7	7(−11)	1.46(6)	9	3(−11)
1650	0	9.27(5)	7	8.5(−11)	4(−11)	9.27(5)	7	7(−11)	7.74(5)	7	7(−11)	9.09(5)	9	3(−11)
1980	0	7.73(5)	7	8.5(−11)	4(−11)	7.73(5)	7	7(−11)	5.55(5)	7	7(−11)	6.10(5)	9	3(−11)
2310	0	6.62(5)	7	1.1(−10)	3(−11)	6.62(5)	7	1(−10)	4.20(5)	7	4(−10)	4.39(5)	9	1(−10)
2640	0	5.80(5)	7	1.1(−10)	3(−11)	5.80(5)	7	1(−10)	3.34(5)	7	4(−10)	3.54(5)	9	1(−10)
2970	0	5.15(5)	7	1.1(−10)	3(−11)	5.15(5)	7	1(−10)	2.74(5)	7	4(−10)	2.68(5)	9	1(−10)
3300	0.01	4.64(5)	7	1.1(−10)	3(−11)	4.64(5)	7	1(−10)	2.29(5)	7	4(−10)	2.20(5)	9	1(−10)
3630	0.01	4.21(5)	7	1.5(−10)	3(−11)	4.21(5)	7	1.5(−10)	1.95(5)	7	3(−10)	1.83(5)	9	6(−10)
3960	0.01	3.54(5)	7	1.5(−10)	3(−11)	3.86(5)	7	1.5(−10)	1.70(5)	7	3(−10)	1.46(5)	9	6(−10)
4290	0.02	3.01(5)	7	1.5(−10)	3(−11)	3.57(5)	7	1.5(−10)	1.49(5)	7	3(−10)	1.22(5)	9	6(−10)
4620	0.02	2.60(5)	7	1.5(−10)	3(−11)	3.31(5)	7	1.5(−10)	1.32(5)	7	3(−10)	1.07(5)	9	6(−10)
4950	0.02	2.26(5)	7	1.3(−10)	5(−12)	3.09(5)	7	5(−11)	1.18(5)	7	3(−10)	9.27(4)	9	4(−10)
5280	0.05	1.99(5)	7	1.3(−10)	5(−12)	2.90(5)	7	5(−11)	1.07(5)	7	3(−10)	8.5(4)	9	4(−10)
5610	0.05	1.76(5)	7	1.3(−10)	5(−12)	2.73(5)	7	5(−11)	9.67(4)	7	3(−10)	7.20(4)	9	4(−10)
5940	0.05	1.57(5)	8	1.3(−10)	5(−12)	2.58(5)	8	5(−11)	8.82(4)	7	3(−10)	6.34(4)	9	4(−10)
6667	0.05	1.27(5)	8	1(−10)	≤ 4(−12)	2.02(5)	8	2.5(−11)	8.10(4)	7	1.5(−10)	5.25(4)	9	4(−10)
7733	0.05	8.50(4)	9	1(−10)	≤ 4(−12)	1.57(5)	9	2.5(−11)	6.77(4)	7	1.5(−10)	4.55(4)	9	4(−10)
8867	0.05	6.50(4)	10	1(−10)	≤ 4(−12)	8.50(4)[f]	10	2.5(−11)	5.42(4)	7	1.5(−10)	3.46(4)	9	4(−10)
10000	0.02	5.20(4)	11	1(−10)	≤ 4(−12)	5.20(4)[f]	11	2.5(−11)	4.41(4)	7	1.5(−10)	2.68(4)	9	4(−10)
13333	0	3.45(4)	12	1(−10)	≤ 4(−12)	3.45(4)[f]	12	2.5(−11)	3.68(4)	7	1.5(−10)	1.70(4)	9	4(−10)

[a] $\rho \propto r^{-1}$ (r < 4000 AU), $\rho \propto r^{-2}$ (r > 4000 AU). This model was used to compute the synthetic spectra in Figs. 1 and 2. Its best N2D+ abundance profile is also listed.

[b] Parameters are held constant inside a given radial shell, except for the rotational speed which is linearly interpolated.

[c] The first three layer densities do not follow a power law but the density profile derived from the ISOCAM absorption map (Pagani et al. 2004)

[d] 2.34(6) reads 2.34×10^6.

[e] Same as note (b) above, but rescaled to maintain a constant total column density of 1.4×10^{23} cm^{-2} (see text).

[f] For the pure $\rho \propto r^{-1}$ case, density is not falling off fast enough to reach ambient cloud density and we correct for this in the last 3 layers.

ment), so that in all cases $\Delta\chi^2 \approx 1$ is equivalent to a 1σ deviation. With this choice, a 3σ variation is equivalent to $\Delta\chi^2_\nu \approx 1.5$ for the first two and $\Delta\chi^2 = 9$ for the last two.

After having run all of the models in our grid, we first used a single, global χ^2_ν combining the four types of measurements to visualize the global fit quality of the various models. Figure 3 plots the global χ^2_ν as a function of temperature, for the various density laws that we investigated. For each density-temperature pair, we plot only the smallest χ^2_ν obtained by adjusting the abundance profile.

It is readily seen that the best fits are found for relatively low overall core temperatures $\simeq 6.5$–8 K. Hence, the (J:1–0) data alone (which dominate the global χ^2_ν) set a strong constraint on the core temperature: indeed, the combination of strong opacity and relatively low intensity in this line requires a low excitation temperature. Because the (J:1–0) line should be thermalized in the dense center, this rules out kinetic temperatures above 8 K (for our adopted total column density of 1.4×10^{23} cm^{-2}).

To further discriminate among models, we then looked individually at the four quality indicators described above, in particular those related to the (J:3-2) lines. Table 3 lists these 4 values, as well as the global χ^2_ν for various models selected from Fig. 3, namely: the best model fits for a given constant core temperature (from 6 to 10 K) and the best model fits for a given density law (with corresponding temperature and abundance distributions given in Table 2).

It can be seen in Table 3 that, for $T_{kin} = 10$ K the (J:3–2) lines are not well-fitted (4 and 5 σ deviations for the CSO and IRAM lines, respectively), hence such a high temperature seems ruled out here. Overall, it is rather difficult to have low χ^2 values in both the (J:1-0) and the (J:3-2) indicators. Only one density-temperature combination (indicated by bold face in Table 3) reaches low values in all 4 individual χ^2_ν and χ^2, each within 1σ of their minimum value, among all the models in Table 3. This combination is referred to as our "best model" in the following, and its inferred physical parameters are discussed in the next section. It can also be noticed in Table 3 that the model with $T_{kin} = 7$ K has a slightly better global χ^2_ν than the best model (2.1 compared to 2.3) but it needs to be eliminated as the (J:3–2) lines are 2 and 3.5 σ off the mark.

3.4. Physical conditions in the L183 main core

3.4.1. Density and temperature structure

Our best density-temperature combination (Fig. 3) has a density law of r^{-1} out to $r \approx 4000$ A.U., consistent with the ISOCAM profile (Pagani et al. 2004) and a slope of r^{-2} beyond. The temperature is constant at 7 K out to 5600 A.U. and increases up to 12 K at the core edge (see Table 2). Increasing the number of warmer layers further failed, as did introducing warmer layers in

Table 3. Quality of fit for selected models from Fig. 3 (see text).

model parameters		(J:1–0) Total area χ^2_ν	5 channels area χ^2_ν	(J:3–2) CSO Total area χ^2	IRAM Total area χ^2	Global χ^2_ν
Best fit for each temperature (adjusted in abundance)						
T = 6 K	$(\rho \propto r^{-1}, r^{-2})$	4.7	3.8	1.5	1.0	3.9
T = 7 K	$(\rho \propto r^{-1}, r^{-2})$	1.9	2.4	12.6	4.1	2.1
T = 8 K	$(\rho \propto r^{-1.5})$	3.1	3.2	0.2	3.4	2.9
T = 9 K	$(\rho \propto r^{-2})$	4.5	5.6	7.5	6.8	4.7
T = 10 K	$(\rho \propto r^{-2})$	4.8	6.3	16.9	28.7	5.5
Best fit for each density law (adjusted in abundance)						
$\rho \propto \mathbf{r^{-1}, r^{-2}}$	**(T = 7 → 12 K)**	**2.4**	**2.7**	**0.3**	**0.0**	**2.3**
$\rho \propto r^{-1}$	(T = 7 → 12 K)	3.7	3.8	1.2	2.9	3.4
$\rho \propto r^{-1.5}$	(T = 7 K)	3.4	3.2	0.1	0.1	2.9
$\rho \propto r^{-2}$	(T = 9 K)	4.5	5.6	7.5	6.8	4.7
Varying the central N_2H^+ abundance at r < 660 AU in the best model						
$X(N_2H^+) =$	10^{-12}	2.6	2.8	0.3	2.0	2.5
$X(N_2H^+) =$	8×10^{-12}	2.6	2.8	0.3	1.0	2.5
$\mathbf{X(N_2H^+) =}$	$\mathbf{2.4 \times 10^{-11}}$	**2.4**	**2.7**	**0.3**	**0.0**	**2.3**
$X(N_2H^+) =$	8×10^{-11}	2.1	2.4	0.5	9.2	2.1

the center. Keeping the temperature constant at 7 K everywhere significantly worsens the fit of the (J:3–2) lines (see Table 3).

The main uncertainty in our temperature determination stems from the assumed total gas column density through the core. Decreasing/increasing it by a factor 1.4 (the typical uncertainty found by Pagani et al. 2004) changes all volume densities by the same factor. To recover the same N_2H^+ emission, we find that the kinetic temperature in our models must be increased/decreased by $\simeq 1$ K (while the N_2H^+ abundance must be scaled accordingly to keep the same column density). Since the dust temperature in the L183 core is 7.5 ± 0.5 K (Pagani et al. 2004), gas and dust are thermalized inside this core within the uncertainties.

A different result was found by Bergin et al. (2006) in the B68 PSC, where outer layers emitting in CO are consistent with a kinetic temperature of 7–8 K while NH_3 measurements indicate a higher temperature of 10–11 K in the inner 40″. This inward increase in gas temperature was attributed to the lack of efficient CO cooling in the depleted core center, and requires an order of magnitude reduction in the gas-to-dust coupling, possibly due to grain coagulation (Bergin et al. 2006). Such an effect is not seen in L183, despite strong CO depletion within the inner 1′= 6600 AU radius. Our best-fit temperature law is more consistent with the thermo-chemical evolution of slowly contracting prestellar cores with standard gas-dust coupling coefficients, which predicts low gas temperatures $\simeq 6$ K near the core center, increasing to $\simeq 14$ K near the core surface (Lesaffre et al. 2005).

3.4.2. N_2H^+ abundance profile

With our best density-temperature model, several N_2H^+ abundance profiles give equally good fits (within $<1\sigma$ on all 4 quality indicators). From all these acceptable abundance profiles, we determined a median profile (Fig. 4a), with error bars representing the range of acceptable values in each layer (though not just any combination of these values will fit the observations). This median profile, whose values are listed in Table 2, is used to compute the fit displayed in Fig. 1. The total N_2H^+ column density is $1.2 \pm 0.1 \times 10^{13}$ cm^{-2}, comparable to the value reported by Dickens et al. (2000) and a factor of ~2 below that in Crapsi et al. (2005).

The maximum N_2H^+ abundance is $1.5^{+0.4}_{-0.3} \times 10^{-10}$, the same as found by Tafalla et al. (2004) in L1498 and L1517. However, there is a definite drop in abundance in the inner layers. This drop is imposed by the fit to the (J:3–2) IRAM central spectrum (9″ beam), which is quite sensitive to variations in the central abundance of N_2H^+, as illustrated in the last rows of Table 3. In contrast, the (J:1–0) line is optically thick at the core center and the χ^2_ν is thus not sensitive to this parameter (cf. Table 3).

Exploring a broad range of abundance profiles, we find only a limited range of possible N_2H^+ abundances in the inner layer of 6″ radius (r < 660 AU): values above 4.5×10^{-11} always produce (J:3–2) emission that is too strong in the central IRAM beam ($\chi^2 > 1$), while abundances below 10^{-12} give a (J:3–2) line that is marginally too weak. We thus derive a range of $2.4 ^{+2.1}_{-2.3} \times 10^{-11}$ for the central N_2H^+ abundance. However, we note that our results for N_2D^+ would favor a value of at least 10^{-11} (see next section).

Compared with the maximum abundance, this gives a volume depletion factor of $\simeq 6^{+13}_{-3}$ at the core center. The median abundance drop is less than (but marginally consistent with) that expected from simple geometrical arguments in PPABC, but it confirms that the leveling of N_2H^+ intensity seen across the dust peak is not due to pure opacity effects. As can be seen from Fig. 4a-b, depletion starts at a density of $5-7 \times 10^5$ cm^{-3} and increases as density goes up, in agreement with the conclusions of PPABC. The abundance also drops slightly in the outermost regions, possibly due to partial destruction by CO. Indeed, the outermost layers of the model have densities of a few 10^4 cm^{-3} which is the limit above which CO starts to deplete in this source (PPABC).

3.4.3. N_2D^+ abundance profile

As seen in Fig.2, good fits to the N_2D^+ data may be obtained with the same density and temperature profile as our best model
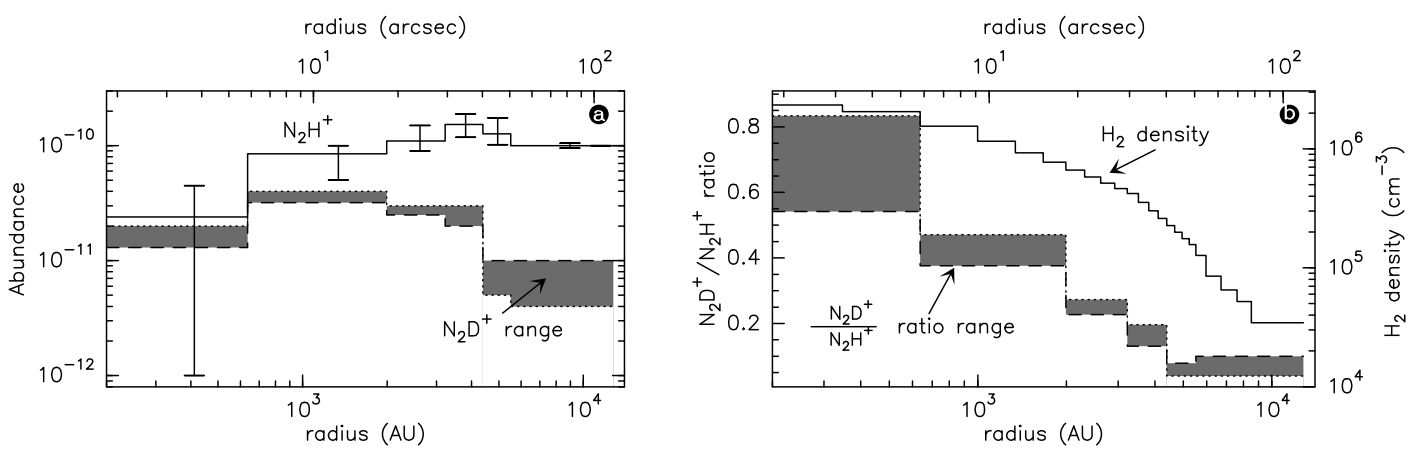
Fig. 4. a N_2D^+ and N_2H^+ abundances for the best model. For N_2H^+, the median abundance profile is plotted (the corresponding model output is displayed in Fig. 1). Error bars indicate the range of possible values (see text). For N_2D^+, a range of values is given: The upper, dotted histogram gives the best fit to the (J:2-1) and (J:3-2) lines and was used to compute the model in Fig. 2. The dashed histogram gives the best fit to the (J:1-0) and (J:2-1) lines. **b** density profile of the best model and N_2D^+/N_2H^+ ratio range

for N_2H^+, although it is difficult to reproduce simultaneously the intensities of both (J:1-0) and (J:3-2) lines. We plot in Fig. 4a the N_2D^+ abundance profiles giving the best fit either for (J:1-0) and (J:2-1) (dashed histogram) or for (J:2-1) and (J:3-2) (dotted histogram). The latter was used to produce Fig. 2. A better fit could be obtained with a temperature of 8 K instead of 7 K. This might indicate that the collisional coefficients with He are somewhat inaccurate to represent those with H_2. Indeed, it has been shown for NH_3 that collisional coefficients with He could differ by a factor up to 4 with respect to those computed with para-H_2 (Willey et al. 2002). A similar problem may occur here (see also the discussion in Daniel et al. 2006a). Therefore we consider that the temperatures are compatible within the uncertainties on the collisional coefficients.

The N_2D^+ abundance profile is quite different from that of N_2H^+ (cf. Fig. 4a). Its abundance is essentially an upper limit beyond 6000 A.U. It increases sharply by about an order of magnitude in the region between 600 and 4000 A.U., then slightly drops by a factor of 2–2.5 in the core center, reaching an abundance between 1.3 and 2×10^{-11}. The low optical depth of the line ($\tau = 0.84$ for the strongest hyperfine component, $J_{FF'}:1_{23}-0_{12}$) allows to measure the contribution of all layers to the emission and to determine with relatively little uncertainty the abundance profile. In particular, for the chosen density and temperature profiles, we find that the abundance of N_2D^+ in the center of the cloud cannot be below 10^{-11}. Hence, we can set tighter constraints on the N_2D^+ abundance at the core center than was possible for N_2H^+.

The N_2D^+/N_2H^+ ratio obtained by comparing the N_2D^+ range of abundances to the median N_2H^+ abundance profile is plotted in Fig. 4b. The deuteration ratio varies from an upper limit ≤ 0.05–0.1 away from the core to a very high factor of 0.7 ± 0.12 in the depletion region. This is larger than what Tine et al. (2000) reported 48″ further north in the same source, and 3–4 times larger than what Crapsi et al. (2005) report towards the PSC. Deuterium enrichment is thus very strong and can certainly be linked to the strong and extended H_2D^+ line detected towards this source (Vastel et al. 2006). If we consider the full range of possible values for N_2H^+ itself, the deuteration ratio in the inner layer ranges from ∼0.3 to ≥ 20. However, as D_2H^+ has not been detected in this source despite its strong H_2D^+ line (Vastel et al. 2006), an enrichment above 1 seems unprobable, suggesting that the central abundance of N_2H^+ is probably at least 10^{-11}.

3.5. Comparison with NH_3 temperature estimates

The low kinetic temperatures ≤ 8 K inferred from our N_2H^+ and N_2D^+ modeling are somewhat lower than previous temperature estimates in L183 from NH_3 inversion lines, which gave values in the range 9–10 K (Ungerechts et al. 1980, Dickens et al. 2000) up to 12 K (Swade et al. 1989). However, these NH_3 spectra were not obtained towards the PSC center itself, and had relatively low angular resolution for the last two. We thus briefly reconsider this issue using our NH_3 spectra obtained at the PSC center position, which also benefit from the higher resolution and beam coupling of the new GBT. The NH_3 spectra were analyzed in the standard way, as discussed by Ho & Townes (1983) and Walmsley & Ungerechts (1983). Using the CLASS NH3 fitting procedure, we found a total opacity of 24 for the (1,1) inversion line and an excitation temperature of 5.5 K, assuming a beam filling-factor of 1 (Figs. 5 & 6). Making the usual assumption of constant temperature on the line of sight and of negligible population in the non-metastable levels, the intensity ratio of the (2,2) to (1,1) main lines indicates a rotation temperature of 8.4 ±0.3 K which should correspond to a kinetic temperature of 8.6 ±0.3 K. Hence, NH_3 emission towards the PSC center indicates an only slightly higher kinetic temperature than N_2H^+ (8.6 ±0.3 K instead of 7 ±1 K), almost equal to that obtained with N_2D^+ (∼8 K). The discrepancy between NH_3 and N_2H^+ is thus not as large as originally thought. Both tracers point to very cold gas in the core, close to thermal equilibrium with the dust.

Reasons for a possible difference between NH_3 and N_2H^+ temperature determinations in L183 include the following: 1) the higher sensitivity of NH_3 to the warmer outer layers of the core, since NH_3 inversion lines are much easier to thermalize ($n_{crit} \approx 2000$ cm^{-3}) than N_2H^+ lines; 2) a slight overestimate in the total column density towards the PSC (reducing it to 10^{23} cm^{-2}, the temperature has to be raised to 8 K to compensate for the density decrease and recover the same N_2H^+ emission); 3) systematic errors introduced by the use of collisional coefficients with He instead of H_2 for N_2H^+ (see Sect. 3.4.3); 4) concerning NH_3, standard hypotheses leading to a puzzling discrepancy between $T_{ex} = 5.5$ K and $T_{rot} = 8.4$ K towards the L183 PSC center, where NH_3 inversion lines should be fully thermalized. Beam dilution has been invoked for giant molecular clouds (thus increasing T_{ex}), but it seems unreasonable to extend this to dense cloud cores (Swade 1989, and references therein). A Monte-Carlo code
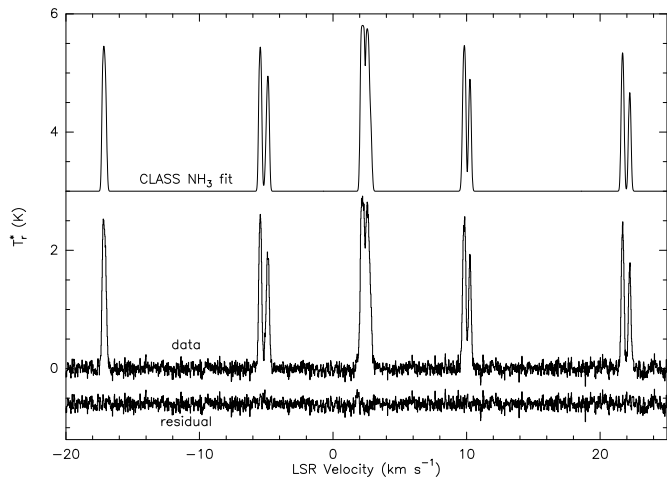
Fig. 5. GBT NH₃ (1,1) inversion line towards the reference position with CLASS NH3 fit (shifted by +3 K for clarity) and fit residual (shifted by -0.6 K). The fit yields a total opacity of 24.2 (±0.4) and a global velocity of 2.3672 (±0.0002) km s^{-1}
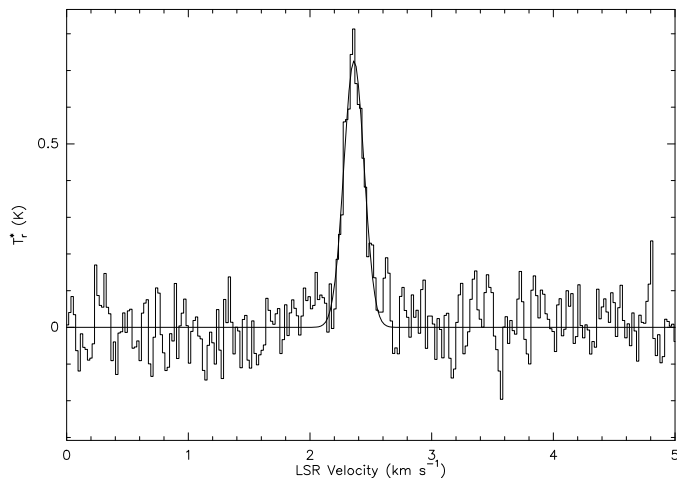

Fig. 6. GBT NH₃ (2,2) inversion line towards the reference position. A Gaussian fit to the main line indicates a velocity of 2.362 (±0.005) km s^{-1}

(or equivalent) would be needed to model NH₃ taking into account both the strong density gradients present in PSCs and the possible population of non-metastable levels at very high densities.

4. Conclusions

We have presented a new Monte-Carlo code (available upon request to the author) to compute more realistically the NLTE emission of N_2H^+ and N_2D^+, taking both line overlap and hyperfine structure into account. This code may be used to infer valuable information on physical conditions in PSCs. Applied to the L183 PSC, we found the following results:

1. The best kinetic temperature to explain N_2H^+ observations is 7±1 K (and ∼ 8 K for N_2D^+) inside 5600 AU, therefore gas appears thermalized with dust in this source.
2. There is no major discrepancy between N_2H^+ and NH₃ measurements. The latter also indicate very cold gas (8.6 ±0.3 K) towards the PSC.
3. We find a noticeable depletion of N_2H^+ by a factor of 6^{+13}_{-3}, and of N_2D^+ by a smaller factor of 2–2.5. This smaller deple-

tion is probably due to a strong (0.7±0.12) deuterium fractionation, consistent with the detection of H_2D^+ in this core.
4. N_2D^+ should be a useful probe of the innermost core regions, thanks to its low optical depth combined with its strong enhancement.

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